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## Effect of different stellar galactic environments on planetary discs I. The solar neighbourhood and the birth cloud of the sun

Jiménez-Torres, J J; Pichardo, B; Lake, G; Throop, H

**Abstract:** We have computed trajectories, distances and times of closest approaches to the Sun by stars in the solar neighbourhood with known position, radial velocity and proper motions. For this purpose, we have used a full potential model of the Galaxy that reproduces the local  $z$ -force, the Oort constants, the local escape velocity and the rotation curve of the Galaxy. From our sample, we constructed initial conditions, within observational uncertainties, with a Monte Carlo scheme for the 12 most suspicious candidates because of their small tangential motion. We find that the star Gliese 710 will have the closest approach to the Sun, with a distance of approximately 0.34 pc in 1.36 Myr in the future. We show that the effect of a flyby with the characteristics of Gliese 710 on a 100 au test particle disc representing the Solar system is negligible. However, since there is a lack of 6D data for a large percentage of stars in the solar neighbourhood, closer approaches may exist. We calculate parameters of passing stars that would cause notable effects on the solar disc. Regarding the birth cloud of the Sun, we performed experiments to reproduce roughly the observed orbital parameters such as eccentricities and inclinations of the Kuiper belt. It is now known that in Galactic environments, such as stellar formation regions, the stellar densities of new born stars are high enough to produce close encounters within 200 au. Moreover, in these Galactic environments, the velocity dispersion is relatively low, typically  $\sim 1$ -3 km s<sup>-1</sup>. We find that with a velocity dispersion of  $\sim 1$  km s<sup>-1</sup> and an approach distance of about 150 au, typical of these regions, we obtain approximately the eccentricities and inclinations seen in the current Solar system. Simple analytical calculations of stellar encounters effects on the Oort Cloud are presented.

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# Effect of Different Stellar Galactic Environments on Planetary Discs I: The Solar Neighbourhood and the Birth Cloud of the Sun.

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## ABSTRACT

We have computed trajectories, distances and times of closest approaches to the Sun by stars in the Solar neighborhood with known position, radial velocity and proper motions. For this purpose we have used a full potential model of the Galaxy that reproduces the local z-force, the Oort constants, the local escape velocity, and the rotation curve of the Galaxy. From our sample we constructed initial conditions, within observational uncertainties, with a Monte Carlo scheme for the twelve most suspicious candidates because of their small tangential motion. We find that the star Gliese 710 will have the closest approach to the Sun, with a distance of approximately 0.34 pc at 1.36 Myr in the future. We show that the effect of a flyby with the characteristics of Gliese 710 on a 100 AU test particle disk representing the Solar system is negligible. However, since there is a lack of 6D data for a large percentage of stars in the Solar neighborhood, closer approaches may exist. We calculate parameters of passing stars that would cause noticeable effects on the Solar disk. Regarding the birth cloud of the Sun, we performed experiments to reproduce roughly the observed orbital parameters such as eccentricities and inclinations of the Kuiper Belt. It is known now that in Galactic environments, such as stellar formation regions, the stellar densities of new born stars, are high enough to produce close encounters within 200 AU. Moreover, in these Galactic environments, the velocity dispersion is relatively low, typically  $\sigma \sim 1\text{--}3$  km s<sup>-1</sup>. We find that with a velocity dispersion of  $\sim 1$  km s<sup>-1</sup> and an approach distance of about 150 AU, typical of these regions, we obtain approximately the eccentricities and inclinations seen in the current Solar system. Simple analytical calculations of stellar encounters effects on the Oort cloud are presented.

**Key words:** Stars:kinematics; Solar Neighborhood:Stellar Dynamics; Oort Cloud:General; Planetary or Debris Discs:General; Birth Cloud of the Sun, Kuiper Belt: Stellar Encounters; Stellar Formation Regions:General.

## 1 INTRODUCTION

Observations of extrasolar planets (e.g. Schneider 2010) show that planetary orbits in other planetary systems are disordered, showing a wide range of eccentricities. For extrasolar planets with semi-major axes  $a \geq 0.1$  AU, the mean of the eccentricity distribution is  $e \approx 0.3$  and the median is  $e \approx 0.24$  (Adams 2010). Thus, available data indicate that planetary systems discovered the last decade are more dynamically active and disordered than our own (Udry & Santos 2007; Adams 2010).

A study of planetary disc dynamics under the stellar influence of different Galactic environments is presented in a set of papers. In this work we introduce the first two Galactic environments related to the Sun: the Solar neighborhood and the birth cloud of the Sun.

The Solar neighborhood has been defined as the region of space centered on the Sun that is much smaller than the overall size of the Galaxy, and whose contents are known with reasonable completeness (Gilmore 1992; Binney & Tremaine 2007). Now, from a theoretical simple approximation, it seems clear that the probability in the current Solar neighborhood to have an important encounter for the planetary system, with another star (i.e. less than 300 AU

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for a solar mass stellar flyby), is almost negligible. Let us for example, consider a typical stellar density for the Solar neighborhood of approximately  $0.05 M_{\odot} \text{ pc}^3$ , we can calculate the mean free path,  $\lambda$ , for approaches within say 300 AU, as the radius of a cross section,  $\sigma$ , and setting 50 km/s as the typical velocity dispersion,  $\sigma_v$ , in the Solar neighborhood, we find that the time necessary, in the current conditions of velocity dispersion and density of the Solar neighborhood, to see an encounter within 300 AU with the Sun (or between any couple of stars in these dynamical conditions), would be,  $t \approx \lambda/\langle v \rangle \approx 1/\sigma n \sigma_v$ , that corresponds to approximately three Hubble times. On the other hand, García-Sánchez et al. (2001), by comparing Hipparcos observations with the stellar luminosity function for star systems within 50 pc of the Sun, estimate that only about one-fifth of the star systems were detected by Hipparcos, and they correct for that incompleteness in the data obtaining about 12 stellar encounters per Myr within one pc of the Sun.

However, in the case of the Solar system it is clear that a rough approximation, is not good enough. First with Hipparcos, and some knowledge of the local and global Galactic potential we are now able to compute orbits of neighboring stars and determinate at good approximation distances, times, etc. to stars in the Solar neighborhood at every time within a few million years. This will improve enormously in the near future with the advent of large surveys of the Milky Way such as GAIA. For now we are not close to having a complete set of 6-dimensional information (position and velocity) of all stars near the Sun. In this paper we take the nearest stars to the Sun with proper motions, parallaxes and radial velocities and we compute their past and future trajectories as well as their closest approach, distance and time. Thus, for our purposes, the Solar neighborhood is the radius of the sphere that contains all the stars (with 6-dimensional parameters known), whose maximum approach to the Sun was or will be less than 5 pc within 10 Myr to the past or to the future, this is approximately 200 pc.

On the other hand, even when, in the current Solar neighborhood conditions, the probability of close stellar approaches is almost negligible, it is now known that as many as 90% of stars appear to form in clusters or groups with  $10^2$  to  $10^3$  members (Carpenter 2000; Lada & Lada 2003; Clark et al. 2005; Adams 2010). Short-lived radioisotopes in Solar meteorites argue that formation happened near at least one massive star, probably in a large cluster (Goswami & Vanhala 2000; Meyer & Clayton 2000; Hester et al. 2004; Looney et al. 2006; Wadhwa et al. 2007). In their early stages, most stars were in relatively crowded environments. In such environments, close stellar encounters would be frequent and affect the stability of planetary systems around the stars (de la Fuente Marcos & de la Fuente Marcos 1997; Laughlin & Adams 1998; Hurley & Shara 2002; Pfahl & Muterspaugh 2006; Spurzem et al. 2010).

Some work has been devoted to searching for stellar perturbers of the cometary disks and clouds. Mathews (1994) identified close approaches for six stars within the next  $5 \times 10^4$  yr, within a radius of about 5 pc. García-Sánchez et al. (1997) started a search for stars passing close to the Sun using Hipparcos data, assuming a straight line motion model. Subsequently García-Sánchez et al. (1999) continued their search by integrating the motion of the candidate stars and the Sun in a local Galactic potential and in a sim-

ple global Galactic Potential. More recently, Bobylev (2010) presents a study of the closest encounters of stars in the Solar neighborhood with the Sun, using also a simple local approximation to integrate orbits.

Of course, a star passing near the Sun has greater effect on the Oort cloud – an extremely extended structure – than on the planetary disk. However, in the solar system, the Kuiper belt has properties acquired early in its history, and some of these are difficult to explain under the assumption that the solar system has been always isolated, such as the excitation of the eccentricities and inclinations in the classic belt; the mass deficit of the Kuiper belt, and the sharp outer edge of the classical belt at approximately 48 AU. Beyond this boundary, only high-eccentricity objects typical of the scattered disk or of the detached population seem to exist, and finally, the existence of some mysterious large bodies (approximately Pluto size), with extreme eccentricities and perihelia (Sedna). One of the most accepted theories to explain the characteristics of Kuiper Belt, is then based on the idea that a close stellar passage could have taken place in the early history of our planetary system (Ida et al. 2000; Kobayashi & Ida 2001). In this paper we experiment with stellar encounters on a 100 AU particle disk, looking for the minimum distance where an encounter is important, and we seek to produce Kuiper belt orbital characteristics such as eccentricities and inclinations.

This paper is organized as follows. In Section 2, we describe the methods and numerical implementation. In Section 3 we present the stellar sample of the Solar neighborhood and our results for this first Galactic environment. In Section 4 we show the results of the Sun’s birth cluster. Conclusions are presented in Section 5.

## 2 METHODOLOGY AND NUMERICAL IMPLEMENTATION

We have implemented two codes to solve the equations of motion. The first calculates stellar trajectories in the Solar vicinity with all their observational orbital parameters known (positions and velocities) in a Milky-Way-like potential. The second simulates debris discs under the influence of stellar encounters.

### 2.1 The Solar Neighborhood Code

We calculate the trajectory of stars with known 6D data (derived from  $\alpha$ ,  $\delta$ , radial velocity and proper motions known) to determinate the distance and time (past or future) of their closest approach to the Sun.

Instead of using a straight line approximation to solve the stellar orbits in the Solar neighborhood, we solved the orbits in a Milky-Way-like Galactic potential (Pichardo et al. 2003, 2004).

The axisymmetric part of the model consists of a background potential with a bulge, a flattened disk (Miyamoto & Nagai 1975) with a scale-height of 250 pc, and a massive halo extending to a radius of 100 kpc. The model has a total mass of  $9 \times 10^{11} M_{\odot}$ , with a local escape velocity of 536 km  $\text{s}^{-1}$ . This potential satisfies observational constraints such as the Galactic rotation curve, with a rotation velocity of 220 km  $\text{s}^{-1}$  at  $R_0 = 8.5$  kpc, the perpendicular force at

the Solar circle, and the Oort constants, among others (for more details see Pichardo et al. 2003, 2004; Martos et al. 2004; Antoja et al. 2009).

Given data in the equatorial system are transformed to Galactic coordinates. The equations of motion are solved with a Bulirsh-Stoer adaptive integrator (Press et al. 1992) that gives relative errors for the integrals of motion (total energy,  $z$  component of angular momentum, for the axisymmetric potential, or Jacobi constant, for the barred or armed potential) of  $10^{-10}$  in the worst case (when the non axisymmetric components of the Galaxy are included). The Sun is located 8.5 kpc from the Galactic nucleus and 0.035 kpc above the plane of the disc. The velocity along the  $x$  axis is  $-9 \text{ km s}^{-1}$  and on the  $y$  axis is  $-220 \text{ km s}^{-1}$ .

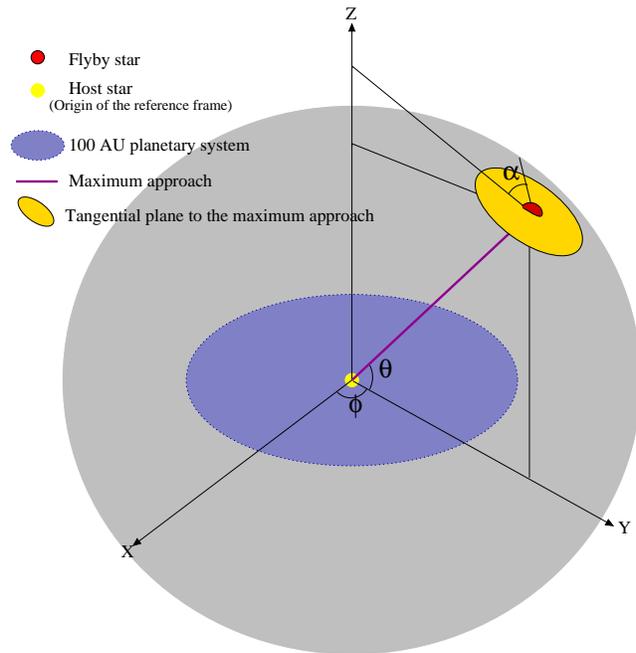
This code computes distances, times, velocities and uncertainties for stars at the moment of maximum approach to the Sun within observational uncertainties. These are obtained from catalogs, and data basis mainly (Simbad, Hipparcos, Nexxus 2) and papers (Bower et al. 2009; Dybczynski 2006; García-Sánchez et al. 1999). Initial conditions are constructed as normal random values within observational uncertainties (note that it is the uncertainty in parallax the one with the normal distribution, not the uncertainty in distance). Ten thousand orbits were computed as initial conditions for each star in order to obtain the final error bars.

In the original version of the code, additionally to the background axisymmetric potential, it included non-axisymmetric features: spiral arms and bar (Pichardo et al. 2003, 2004). In the same way as García-Sánchez et al. 2001, who constructed a potential including a very simple axisymmetric background and weak spiral arms (pitch angle of  $\sim 6.9^\circ$ ), we conclude that the difference between using a full Galactic model and an axisymmetric model among the Solar neighborhood stars is negligible (within  $\approx 3\%$  for the most distant stars to the Sun in this study). Although there is evidence that the existence of spiral arms and bars have direct influence on the local stellar distribution and evolution (Dehnen 2000; Chakrabarty 2007; Antoja et al. 2009), the time we run the orbits is rather short ( $\leq 10^7$  years), much less than a dynamical time, and not enough to produce a noticeable difference in the orbits produced by the non-axisymmetric potential. To facilitate the calculations we performed the final calculations only with the axisymmetric background potential of the Galaxy, since the bar and arms are not of importance in this time scale.

## 2.2 The Stellar Encounter Code

The second code is a 3D particle code that calculates the interaction of a planetary system (a central star and a disk of test particles) with a second stellar body (a flyby star).

The sampling of test particles goes as  $a \propto n^{-3/2}$ , where  $a$  is the initial radius of a particle orbit and  $n$  is the number of the orbit; these test particles are affected by the gravitational forces of both the central star and the flyby star; and the equations of motion are solved for the central stars (Suns) non-inertial frame of reference. The code calculates the main orbital characteristics of each of the test particles simulating a debris disk after the flyby; it calculates the semi-major axis, the eccentricity, and the inclinations. The Bulirsh-Stoer integrator gives a maximum relative error be-



**Figure 1.** Schematic figure of a stellar encounter on a planetary system by a flyby star. The disc is simulated with test particles. The initial conditions are particles on circular orbits with zero inclination.

fore the flyby of  $10^{-14}$  and  $10^{-13}$  in the energy and angular momentum integrals, respectively.

In Figure 1, we show the schematics of the relevant parameters used on the code of a stellar encounter. The dark disk in the center of the system represents the planetary disk, the gray sphere represents the radius of the minimum distance of the flyby, and the bright disk is tangent to the sphere at the point of minimum distance. The flyby attack angles are:  $\phi$ , the azimuthal angle with respect to the disc, it goes from  $0^\circ$  to  $360^\circ$ ;  $\theta$ , the polar angle with respect to the disc, goes from  $-90^\circ$  to  $90^\circ$ ; and  $\alpha$ , the angle between the flyby plane orbit and the symmetry axis of the planetary disc, it goes from  $0^\circ$  to  $360^\circ$ .

With the purpose of covering several Galactic stellar environments, and to acquire a physical idea of what parameters are the most relevant in this kind of interaction, besides the specific applications of this code to the Sun's birth cluster and the current Solar environment, we computed more than a thousand experiments covering the different Galactic ranges in distance (stellar closest approximation), velocities, angles of flyby trajectories and stellar masses. For distance: 5 to 1000 AU, for velocity: 1 to 180 km/s, for mass: 1 Jupiter mass to 4 Solar masses. Regarding flyby attack angle  $\phi$ , due to the symmetry of the problem, this entrance angle is indistinct; we took for our experiments  $\phi = 0^\circ$ .

In the application presented in this paper the effects of the exact direction of entrance of the flyby resulted of no importance, this, due to the fact that: a) in the current Solar neighborhood stars will not come close enough for the angle to be important and the effect of encounters at such large distances would be negligible, even for objects in the Oort cloud in the majority of cases; b) in the case of the

birth cloud of the Sun, stars approaches might have been important, however in this case the orientation of discs with respect to flybys are unknown and probably rather random.

Since we are talking about one single encounter in this Galactic environment, we chose a general interaction, this is  $45^\circ$  for  $\theta$  and  $\alpha$ , that should produce an intermediate effect produced by all possible inclinations of the orbital plane. This means, the flyby orbit will enter at  $45^\circ$  with respect to the plane of the disc, and  $45^\circ$  with respect to the disc axis. It is worth to mention, that the results are nearly insensitive to changes of  $\alpha$ , while changes in  $\theta$  produce diferent results being larger for  $\theta = 0^\circ$ . A global study of the parameters will be presented on a future paper.

### 3 RESULTS

Here we present the first two Galactic regions studied in this work.

#### 3.1 The Solar Neighborhood

For the Solar neighborhood, we looked for the closest stars to the Sun with known position, radial velocity and proper motions. It is worth to mention that for the nature of the sample, this is, we are able only to compute stars with 6D known information from literature, our sample is biased to the closest, more massive and with the largest proper motion stars. However, this bias actually represents the most important candidates to perturb our planetary disc. Considering this bias, from the more than a thousand stars in the full sample, we find that 67 stars passed or will pass within 3 pc from the Sun, between the last 8 Myr and the next 5 Myr, this is roughly, 5 stars per Myr. Only one of this 67 set, will barely perturb the outer Oort cloud, and the planetary disc will not notice that star in the dynamical sense. On the other hand, it should be noticed that there could be a star of which we do not have 6D information that could come even closer than G710. Even in our sample, we could have stars that will approach to the Sun several times along the stellar trajectories around the Galaxy, but in this case, there is no model of the Galaxy good enough to warranty precise orbits for these stars for more than a few million years.

##### 3.1.1 The Stellar Sample

We used parallax and proper-motion data and uncertainties from the Hipparcos catalog for 1167 nearby stars suspected to pass close to the Solar System. We take radial velocity measurements from different catalogs (Hipparcos, Simbad, Nexxus2) and papers (Bower et al. 2009; Dybczynski 2006; García-Sánchez et al. 1999) to produce six-dimensional trajectories of the whole sample of stars, and calculate the distance at the point of maximum approach to the Sun, the time and velocity (the whole table is available on request via e-mail). From this sample, we produced a subsample of 67 stars with their observational uncertainties and calculated closest approach distances to the Sun, within 3 pc. From this subsample, we ran 34 stars toward the past (those receding now), and 33 stars toward the future (those approaching now) looking for those passages that might perturb on

the Solar System. In particular, these 67 stars are potential perturbers of the Oort Cloud (García-Sánchez, et al. 1999).

The stars with the closest approach distances, within 3 pc from the Sun, are listed in Tables 1 and 2 sorted first by receding versus approaching then by parallax. Columns are arranged as follows, star's name, parallax, radial velocity, minimum approach distance to the Sun, approach time, and relative velocity. These predicted passages are contained in a time interval of about -8 and +4 Myr. The close-approach distances vs. time are shown in Figure 2. In the upper panel of this figure we show the variation with time of the separation distance between each star and the Sun, at the moment of maximum approach, for all stars in our sample (of 1167 stars) which closest distance is less than 30 pc. The lower frame is a zoom that shows the stars with the closest approach distances, within 3 pc (Tables 1 and 2), both plots for the time interval, from -8 Myr to 4 Myr. Stars coming within 2-3 pc may perturb the Oort cloud. The star with the closest future approach is Gliese 710 (GJ 710). The predicted minimum distance for this star is 0.34 pc, in 1.36 Myr. This is the only star in the sample with an approach distance less than the radius of the Oort cloud  $\sim 0.48$  pc.

##### 3.1.2 The Most Interesting Candidates

From this sample, if we ignored the uncertainties, we could conclude that none of these stars would approach to the Sun closer than  $\sim 70,000$  AU (more than 200 times beyond than necessary to start affecting the outskirts of our planetary system). We have considered however, observational uncertainties in this work. The stars that may produce important (even dangerous) approaches to the Sun, are those whose tangential velocity were potentially zero (simultaneously in  $\alpha$  and  $\delta$ ). We will then check on those stars with a tangential velocity (or approach distance to the Sun), is consistent with zero at  $3\sigma$ . This includes both stars with small tangential velocity and those with uncertain tangential velocities. This makes a total number of 12, 8 to the past and 4 to the future, from the whole sample of 67 objects.

For these 12 stars (labeled with an asterisc in Tables 1 and 2), we have separately produced in Figures 3 and 4, histograms, eight to the past and four to the future. In each histogram we computed 10,000 random realizations with a normal distribution for the initial conditions within all the observational uncertainties, with the notable exception of the parallax, where the value was restricted to positive values. In these diagrams we analyze the most likely approach distance sun-star toward the past and future respectively. All diagrams were done with the same bin size of 0.01 pc and the same axes (for all stars in the corresponding diagram) for better see differences between the stellar approaches to the past or to the future.

The number of approaches closer than 0.01 pc is of 3 to the past and 8 to the future, this is equivalent to a probability of 0.03% to the past and 0.08% to the future of an approach to less than 0.01 pc. A more interesting distance is of course one that could cause a noticeable effect on the planetary disc, that we know now is approximately of 300 AU, as we will show in Section 3.1.4. The probability of a closer than 300 AU approach is then of  $\sim 0.0017\%$  from the 8-stars to the past and of  $\sim 0.0006\%$  from the 4-stars to the future. The probabilities with these stars to pass close

**Table 1.** Astronomical data for the 34 stellar subsample run toward the past.

Star Name	Parallax (arcsec)	PM( $\alpha$ ) (arcsec/yr)	PM( $\delta$ ) (arcsec/yr)	$V_r$ (km/s)	Miss Distance (pc)	Time ( $10^3$ yr)	Vel $_{App}$ (km/s)
GJ65B(LHS10)	0.3737±0.0060	3.3210±0.0050	0.5620±0.0050	29.0±2.0	2.21±0.07	-28.5±0.9	51.64±1.38
H16537(GJ144)	0.3108±0.0009	-0.9764±0.0010	0.0180±0.0009	15.5±0.9	2.23±0.08	-105.5±0.4	21.50±0.73
H5643(GJ54.1)	0.2691±0.0076	1.2101±0.0052	0.6470±0.0039	28.0±5.0	2.43±0.31	-74.4 <sup>+3.7</sup> <sub>-0.6</sub>	36.99±4.22
H24186(GJ191)	0.2553±0.0009	6.5061±0.0010	-5.7314±0.0009	245.5±2.0	2.15±0.02	-10.9±0.1	293.58±1.90
H30920(GJ234A)	0.2429±0.0026	0.6947±0.0030	-0.6186±0.0025	24.0±5.0	2.48±0.38	-106.7 <sup>+8.4</sup> <sub>-2.8</sub>	30.09±4.41
H103039(LP816-60)	0.1820±0.0037	-0.3067±0.0038	0.0308±0.0041	15.8±0.6	2.49±0.14	-270.2±7.8	17.72±0.61
H33226(GJ251)	0.1813±0.0019	-0.7293±0.0021	-0.3993±0.0013	36.0±10.0	2.85±0.71	-109.8 <sup>+15.7</sup> <sub>-8.8</sub>	42.06±9.35
H40501(GJ2066)	0.1090±0.0018	-0.3750±0.0022	0.0601±0.0015	62.2±0.1	2.35±0.08	-134.7±2.1	64.36±0.13
H14754(HD20523)	0.0985±0.0015	0.0419±0.0015	-0.1041±0.0019	65.9±0.1	0.83±0.03	-149.6±2.4	66.12±0.11
H22738(GJ2036A)	0.0890±0.0036	0.1329±0.0043	0.0739±0.0038	40.1±10.0	2.22±0.90	-263.2 <sup>+52.4</sup> <sub>-112.5</sub>	40.92±10.95
H26335(GJ208)	0.0879±0.0013	-0.0026±0.0013	-0.0576±0.0009	22.7±5.0	1.55±0.52	-481.1 <sup>+90.0</sup> <sub>-188.3</sub>	22.91±5.56
H95326(CCDMJ19236)	0.0780±0.0579	0.1017±0.0749	-0.0216±0.0417	35.6±0.4	2.24±26.40 <sup>a</sup>	-341.3 <sup>+123.5</sup> <sub>-299.0</sub>	36.16 <sup>+22.22</sup> <sub>-0.62</sub>
H27887(GJ2046)	0.0780±0.0005	-0.0518±0.0006	-0.0604±0.0006	30.4±0.2	2.02±0.04	-402.1±3.9	30.79±0.22
H8709(GJ3121)	0.0630±0.0038	0.0012±0.0043	0.1241±0.0024	64.0±3.0	2.29±0.33	-237.4±19.5	64.68±3.34
H12351(GJ1049)	0.0610±0.0013	-0.0190±0.0013	0.0303±0.0013	26.2±10.0	1.74±3.16 <sup>a</sup>	-604.5 <sup>+188.0</sup> <sub>-640.4</sub>	26.37±10.96
GJ54.2B(HD7438)	0.0574±0.0009	0.1320±0.0022	0.2830±0.0021	16.0±5.0	14.80±1.38	-296.5 <sup>+78.7</sup> <sub>-24.5</sub>	30.36 <sup>+3.53</sup> <sub>-2.23</sub>
H32362(GJ242)	0.0570±0.0008	-0.1152±0.0007	-0.1909±0.0005	211.1±0.9	1.54±0.05	-80.6±1.2	211.91±1.01
GJ401B(LHS290)	0.0569±0.0252	-1.8660±0.0249	-0.6610±0.0184	40.2±6.5	17.08 <sup>+251.76</sup> <sub>-5.07</sub>	-24.0±10.4	169.76 <sup>+2511.15</sup> <sub>-43.50</sub>
H27288(GJ217.1)	0.0460±0.0007	-0.0148±0.0006	-0.0012±0.0005	25.6±5.0	1.31±0.38	-827.1 <sup>+142.6</sup> <sub>-287.0</sub>	25.65±5.61
H13772(HD18455)	0.0445±0.0026	0.0302±0.0024	-0.0369±0.0018	50.4±0.2	2.25±0.30	-431.4±26.6	50.69±0.23
H26373(HD37572)	0.0419±0.0017	0.0253±0.0023	0.0000±0.0022	32.4±0.2	2.08±0.25	-714.4±31.1	32.55±0.23
T100111(HD351880)	0.0400±0.0327	-0.0100±0.0340	0.0074±0.0270	26.1±0.3	1.40±70.48 <sup>a</sup>	-933.4 <sup>+370.5</sup> <sub>-1163.6</sub>	26.15 <sup>+20.85</sup> <sub>-0.34</sub>
H13769(GJ120.1C)	0.0389±0.0015	0.0154±0.0015	-0.0325±0.0012	49.6±0.5	2.26±0.21	-502.6±21.2	49.83±0.56
H93506(HD176687)	0.0370±0.0014	-0.0141±0.0026	0.0037±0.0020	22.0±5.0	2.25±0.91	-1192.9 <sup>+231.3</sup> <sub>-545.0</sub>	22.08±5.60
H30067(HD43947)	0.0360±0.0009	-0.0175±0.0010	-0.0143±0.0007	40.5±2.0	2.04±0.19	-667.1±41.7	40.60±2.24
S14576(ALGOL)	0.0350±0.0009	0.0024±0.0008	-0.0014±0.0009	4.0±0.9	2.54±1.86 <sup>a</sup>	-6893.3 <sup>+1315.1</sup> <sub>-2911.5</sub>	4.05±1.00
H30344(HD44821)	0.0340±0.0008	-0.0030±0.0005	0.0041±0.0006	14.4±0.2	1.44±0.17	-1990.0±59.0	14.44±0.23
H54806(HD97578)	0.0310±0.0141	-0.0117±0.0095	0.0018±0.0099	23.5±1.0	2.48±15.52 <sup>a</sup>	-1333.5 <sup>+371.4</sup> <sub>-1355.2</sub>	23.59 <sup>+1.64</sup> <sub>-1.10</sub>
H31626(HD260564)	0.0290±0.0021	0.0095±0.0018	-0.0332±0.0013	82.7±5.0	2.35±0.44	-405.8 <sup>+35.4</sup> <sub>-50.0</sub>	82.89±5.61
H26624(HD37594)	0.0240±0.0007	-0.0039±0.0006	0.0021±0.0004	22.4±1.3	1.67±0.27	-1815.0±134.9	22.43±1.46
T31821(HD47787)	0.0210±0.0021	-0.0343±0.0017	0.0274±0.0025	18.3±0.6	22.66±4.65	-1966.5±128.1	20.84±0.82
H99483(HIP99483)	0.0130±0.0067	0.0013±0.0214	0.0003±0.0251	25.0±0.2	1.95±129.44 <sup>a</sup>	-3001.3±1154.8	25.07 <sup>+13.47</sup> <sub>-0.18</sub>
H40317(HD68814)	0.0120±0.0016	-0.0007±0.0014	0.0020±0.0012	34.2±0.2	2.02±2.15 <sup>a</sup>	-2378.2 <sup>+275.3</sup> <sub>-428.3</sub>	34.27±0.27
H101573(HIP101573)	0.0050±0.0023	-0.0002±0.0024	0.0004±0.0022	43.7±0.5	2.55±21.96 <sup>a</sup>	-4443.3±1418.0	44.15 <sup>+0.91</sup> <sub>-0.59</sub>

<sup>a</sup> The uncertainty value marked here corresponds to the radius of a cylinder where the locus of the closest approach is likely to lie.

enough to affect the planetary disc of the Solar system are extremely low, even for the star that will approach closest (GJ 710).

### 3.1.3 Comparison between the Straight Line Approximation, a Local, and Global Potentials

Following García-Sánchez et al. (2001), we compare stellar orbits with our global potential to those using the straight line approximation and a local potential approximation from

their paper. For distances of the stars larger than 50 pc the local and straight line approximations give important differences with respect to a global model, even with times as short as 1 Myr.

García-Sánchez et al. (2001) concluded that within a time interval of  $\pm 10$  Myr from the present time, the predicted encounters are fairly well determined for most of the candidate stars. They are not altered significantly by the use of alternative Galactic potential models or by varying the plausible values of the Galactic parameters. The most inter-

**Table 2.** Astronomical data for the 33 stellar subsample run toward the future.

Star Name	Parallax (arcsec)	PM( $\alpha$ ) (arcsec/yr)	PM( $\delta$ ) (arcsec/yr)	$V_r$ (km/s)	Miss Distance (pc)	Time ( $10^3$ yr)	Vel $_{App}$ (km/s)
H70890(PROXIMA)	0.7723 $\pm$ 0.0024	-3.7756 $\pm$ 0.0015	0.7682 $\pm$ 0.0018	-21.7 $\pm$ 0.5	0.95 $\pm$ 0.01	26.7 $\pm$ 0.1	32.10 $\pm$ 0.39
H71683(AlphCenA)	0.7421 $\pm$ 0.0014	-3.6200 $\pm$ 0.0015	0.7100 $\pm$ 0.0012	-20.7 $\pm$ 0.9	1.01 $\pm$ 0.02	27.7 $^{+0.1}_{-0.2}$	31.37 $\pm$ 0.67
H71681(AlphCenB)	0.7421 $\pm$ 0.0014	-3.6004 $\pm$ 0.0261	0.9521 $\pm$ 0.0198	-24.6 $\pm$ 0.9	0.94 $\pm$ 0.02	27.7 $\pm$ 0.2	34.22 $\pm$ 0.74
H87937(BARNARD)	0.5490 $\pm$ 0.0016	-0.7978 $\pm$ 0.0016	10.3269 $\pm$ 0.0013	-106.8 $\pm$ 0.2	1.17 $\pm$ 0.01	9.8 $\pm$ 0.1	139.30 $\pm$ 0.25
H54035(GJ411)	0.3924 $\pm$ 0.0009	-0.5802 $\pm$ 0.0008	-4.7671 $\pm$ 0.0008	-85.0 $\pm$ 0.9	1.44 $\pm$ 0.01	20.0 $\pm$ 0.1	102.91 $\pm$ 0.84
S32349(SIRIUS)	0.3792 $\pm$ 0.0016	-0.5460 $\pm$ 0.0013	-1.2231 $\pm$ 0.0012	-9.4 $\pm$ 0.9	2.30 $\pm$ 0.06	65.8 $^{+3.0}_{-4.5}$	19.20 $\pm$ 0.50
H92403(GJ729)	0.3365 $\pm$ 0.0018	0.6376 $\pm$ 0.0022	-0.1925 $\pm$ 0.0015	-4.0 $\pm$ 2.0	2.73 $\pm$ 0.22	111.8 $^{+24.6}_{-57.9}$	10.20 $^{+1.12}_{-0.55}$
GJ905(LHS549)	0.3160 $\pm$ 0.0020	0.0850 $\pm$ 0.0050	-1.6150 $\pm$ 0.0050	-81.0 $\pm$ 5.0	0.91 $\pm$ 0.06	35.1 $\pm$ 2.1	84.56 $\pm$ 5.39
H57548(GJ447)	0.2996 $\pm$ 0.0022	0.6056 $\pm$ 0.0021	-1.2192 $\pm$ 0.0019	-31.0 $\pm$ 0.2	1.91 $\pm$ 0.03	71.0 $\pm$ 0.3	37.75 $\pm$ 0.21
GJ866A(LHS68)	0.2895 $\pm$ 0.0050	2.3640 $\pm$ 0.0050	2.2360 $\pm$ 0.0050	-60.0 $\pm$ 2.0	2.29 $\pm$ 0.08	31.5 $^{+0.1}_{-0.2}$	80.24 $\pm$ 1.82
H104214(GJ820A)	0.2871 $\pm$ 0.0015	4.1551 $\pm$ 0.0010	3.2589 $\pm$ 0.0012	-64.3 $\pm$ 0.9	2.80 $\pm$ 0.03	18.7 $\pm$ 0.1	108.34 $\pm$ 0.73
H110893(GJ860A)	0.2495 $\pm$ 0.0030	-0.8702 $\pm$ 0.0030	-0.4711 $\pm$ 0.0030	-24.0 $\pm$ 5.0	2.47 $\pm$ 0.36	101.2 $^{+2.2}_{-7.2}$	30.49 $\pm$ 4.37
H85605(CCDMJ)	0.2027 $\pm$ 0.0395	0.0973 $\pm$ 0.0267	0.3489 $\pm$ 0.0413	-21.1 $\pm$ 0.2	1.84 $^{+1.20}_{-0.47}$	197.0 $^{+36.5}_{-25.4}$	22.74 $^{+1.18}_{-0.52}$
H86214(GJ682)	0.1983 $\pm$ 0.0024	-0.7101 $\pm$ 0.0028	-0.9380 $\pm$ 0.0021	-60.0 $\pm$ 10.0	2.14 $\pm$ 0.36	67.4 $\pm$ 7.8	66.27 $\pm$ 10.12
H97649(GJ768)	0.1944 $\pm$ 0.0009	0.5368 $\pm$ 0.0007	0.3855 $\pm$ 0.0007	-26.1 $\pm$ 0.9	2.70 $\pm$ 0.08	139.5 $\pm$ 2.4	30.68 $\pm$ 0.86
H57544(GJ445)	0.1855 $\pm$ 0.0014	0.7432 $\pm$ 0.0016	0.4804 $\pm$ 0.0012	-119.0 $\pm$ 5.0	1.01 $\pm$ 0.05	43.0 $\pm$ 1.9	121.13 $\pm$ 5.53
H86990(GJ693)	0.1721 $\pm$ 0.0022	-1.1199 $\pm$ 0.0021	-1.3525 $\pm$ 0.0015	-115.0 $\pm$ 21.0	2.25 $\pm$ 0.44	42.0 $^{+7.0}_{-4.9}$	124.76 $\pm$ 21.61
H99461(GJ783A)	0.1652 $\pm$ 0.0009	0.4569 $\pm$ 0.0009	-1.5749 $\pm$ 0.0006	-129.8 $\pm$ 0.9	2.06 $\pm$ 0.03	40.3 $\pm$ 0.3	138.07 $\pm$ 0.96
H86961(GJ2130A)	0.1618 $\pm$ 0.0113	-0.0498 $\pm$ 0.0929	-0.3198 $\pm$ 0.0473	-28.9 $\pm$ 0.9	1.93 $\pm$ 0.41	188.8 $\pm$ 12.4	30.42 $\pm$ 1.09
H86963(GJ2130B)	0.1618 $\pm$ 0.0113	-0.0776 $\pm$ 0.0136	-0.2701 $\pm$ 0.0078	-27.4 $\pm$ 0.9	1.78 $\pm$ 0.28	202.3 $\pm$ 13.7	28.61 $\pm$ 0.97
H83945(GJ3991)	0.1378 $\pm$ 0.0090	0.3339 $\pm$ 0.0081	-0.2780 $\pm$ 0.0103	-45.0 $\pm$ 10.0	2.29 $\pm$ 0.71	142.0 $^{+37.8}_{-22.1}$	47.42 $\pm$ 10.58
H93449(RCrA)	0.1218 $\pm$ 0.0682	-0.0344 $\pm$ 0.0975	0.0506 $\pm$ 0.0520	-36.0 $\pm$ 5.0	0.54 $\pm$ 27.21 <sup>a</sup>	222.0 $^{+462.7}_{-76.4}$	36.08 $^{+14.33}_{-5.55}$
H77257(GJ598)	0.0851 $\pm$ 0.0008	-0.2255 $\pm$ 0.0008	-0.0685 $\pm$ 0.0007	-66.4 $\pm$ 0.9	2.28 $\pm$ 0.06	166.5 $\pm$ 2.8	67.69 $\pm$ 1.00
H116727(GJ903)	0.0725 $\pm$ 0.0005	-0.0489 $\pm$ 0.0005	0.1272 $\pm$ 0.0004	-42.4 $\pm$ 0.9	2.84 $\pm$ 0.08	305.0 $\pm$ 6.9	43.33 $\pm$ 0.99
H6379(GJ56.5)	0.0595 $\pm$ 0.0006	-0.0341 $\pm$ 0.0005	-0.0345 $\pm$ 0.0006	-22.7 $\pm$ 2.0	2.82 $\pm$ 0.29	703.4 $^{+79.7}_{-55.7}$	23.03 $\pm$ 2.22
H89825(GJ710)	0.0518 $\pm$ 0.0014	0.0017 $\pm$ 0.0014	0.0021 $\pm$ 0.0011	-13.9 $\pm$ 2.0	0.34 $\pm$ 0.28 <sup>a</sup>	1357.8 $^{+312.1}_{-172.7}$	13.90 $\pm$ 2.25
H113421(HD217107)	0.0507 $\pm$ 0.0008	-0.0061 $\pm$ 0.0008	-0.0160 $\pm$ 0.0006	-14.0 $\pm$ 0.6	2.22 $\pm$ 0.16	1355.5 $\pm$ 67.0	14.16 $\pm$ 0.67
H38228(HD63433)	0.0458 $\pm$ 0.0009	-0.0093 $\pm$ 0.0010	-0.0118 $\pm$ 0.0007	-16.5 $\pm$ 0.2	2.08 $\pm$ 0.16	1281.6 $\pm$ 31.0	16.58 $\pm$ 0.22
H105766(GJ4194)	0.0389 $\pm$ 0.0006	0.0412 $\pm$ 0.0006	0.0396 $\pm$ 0.0006	-76.9 $\pm$ 0.2	2.32 $\pm$ 0.07	324.2 $\pm$ 5.3	77.22 $\pm$ 0.23
H20359(GJ168)	0.0326 $\pm$ 0.0020	-0.0348 $\pm$ 0.0018	0.0111 $\pm$ 0.0015	-78.5 $\pm$ 5.0	2.07 $\pm$ 0.34	380.3 $\pm$ 37.0	78.68 $\pm$ 5.61
H21386(HD26367)	0.0273 $\pm$ 0.0015	0.0135 $\pm$ 0.0013	0.0090 $\pm$ 0.0017	-50.7 $\pm$ 2.0	2.01 $\pm$ 0.29	704.1 $\pm$ 51.3	50.80 $\pm$ 2.25
H85661(HD158576)	0.0115 $\pm$ 0.0008	0.0010 $\pm$ 0.0009	0.0000 $\pm$ 0.0004	-46.0 $\pm$ 1.7	0.97 $\pm$ 0.89 <sup>a</sup>	1847.7 $\pm$ 157.1	46.02 $\pm$ 1.89
H94512(HD1779939)	0.0085 $\pm$ 0.0009	-0.0001 $\pm$ 0.0007	-0.0005 $\pm$ 0.0005	-30.1 $\pm$ 2.0	1.49 $\pm$ 1.95 <sup>a</sup>	3821.3 $\pm$ 480.8	30.10 $\pm$ 2.20

<sup>a</sup> The uncertainty value marked here corresponds to the radius of a cylinder where the locus of the closest approach is likely to lie.

esting result is the future passage of Gliese 710 through the outer Oort cloud. This result is in good agreement with the predictions using other Galactic potential models, the prediction of this stellar passage is not model dependent owing to its proximity to the Sun. They integrated the trajectories using three different models of the Galactic potential: a local potential model, a global potential model and a perturbative potential model. The agreement between their models was generally good.

We are interested in knowing the approximate distance to the stars, where the straight line and local approximations differ significantly from the global model. To do so, we use

the data in Table 2 of García-Sánchez et al. (1999) for the closest encounters, using a simple rectilinear motion of the stars, and also the data for the closest encounters using a local Galactic potential.

In Table 5: we show 142 stars close to the Sun (within approximately 190 pc). From this sample, we obtain 74 ( $\sim 52\%$ ) stars for which the difference among the approximations to the potential employed (global, local or linear) is negligible. In Figure 5, we show the relative errors of the closest approaches (missing distance errors), between the global *vs.* straight line approximations (blue triangles), and between the global *vs.* local approximations (red squares),

**Table 3.** Relative errors between different approximations to the Galactic potential. Column 1 is the name of the object; Columns 2, 3 and 4 show the close-approach distance to the Sun in pc with the global (from this work), local and straight line approximations (García-Sánchez et al. 2001,1999); Columns 5,6 and 7 present the relative errors (G=Global, Lo=Local, Li=Linear) and the last column is the current distance to the star in pc.

Star Name	Global	Local Missing Distance (pc)	Linear	(G-Lo)/G	(G-Li)/G (%)	(Lo-Li)/Lo	Dist (pc)
GJ 710	0.337	0.336	0.343	0.3	1.8	2.1	19.30
HD 158576	0.938	0.846	0.753	9.8	19.7	11.0	86.81
PROXIMA	0.954	0.954	0.954	0.0	0.0	0.0	1.29
ALPHA CENTAURI A	0.973	0.973	0.973	0.0	0.0	0.0	1.35
ALPHA CENTAURI B	0.975	0.975	0.975	0.0	0.0	0.0	1.35
AC + 79 3888	1.007	1.007	1.007	0.0	0.0	0.0	5.39
GJ 620.1B	4.259	1.139	1.139	73.3	73.3	0.0	12.81
BARNARD STAR	1.144	1.143	1.143	0.1	0.1	0.0	1.82
HD 351880	1.434	1.439	1.445	0.3	0.8	0.4	25.27
LALANDE 21185	1.440	1.440	1.440	0.0	0.0	0.0	2.55
SAO 75395	2.469	1.448	2.688	41.4	8.9	85.6	118.34
HD 179939	1.444	1.451	1.025	0.5	29.0	29.4	117.10
GJ 208	1.600	1.600	1.599	0.0	0.1	0.1	11.38
HD 37594	1.637	1.610	1.598	1.6	2.4	0.7	41.39
GJ 217.1	1.645	1.637	1.629	0.5	1.0	0.5	21.52
HIP 99483	1.797	1.653	1.379	8.0	23.3	16.6	74.13
HD 35317	1.775	1.755	1.735	1.1	2.3	1.1	58.04
GJ 2130 B	1.782	1.782	1.782	0.0	0.0	0.0	6.18
HD 19995	1.254	1.811	2.653	44.4	111.6	46.5	68.54
HIP 101573	2.072	1.821	1.898	12.1	8.4	4.2	187.62
CCDM 17296 + 2439 B	1.837	1.837	1.837	0.0	0.0	0.0	4.93
GJ 358	1.875	1.875	1.875	0.0	0.0	0.0	9.49
ROSS 154	1.881	1.881	1.881	0.0	0.0	0.0	2.97
HD 68814	1.950	1.909	1.990	2.1	2.1	4.2	82.10
ROSS 128	1.911	1.911	1.911	0.0	0.0	0.0	3.34
GJ 2130 A	1.929	1.929	1.929	0.0	0.0	0.0	6.18
GJ 860 A	1.949	1.949	1.949	0.0	0.0	0.0	4.01
HD 33487	2.001	1.977	1.954	1.2	2.3	1.2	41.63
HD 43947	2.015	2.016	2.016	0.0	0.0	0.0	27.53
HD 26367	2.019	2.028	2.038	0.4	0.9	0.5	36.66
GJ 271 A	2.044	2.038	2.029	0.3	0.7	0.4	18.03
GJ 168	2.074	2.074	2.075	0.0	0.0	0.0	30.69
GJ 144	2.135	2.135	2.135	0.0	0.0	0.0	3.22
HD 63433	2.150	2.138	2.121	0.6	1.3	0.8	21.82
GJ 682	2.140	2.140	2.140	0.0	0.0	0.0	5.04
GJ 120.1	2.243	2.245	2.246	0.1	0.1	0.0	22.48
GJ 693	2.253	2.253	2.253	0.0	0.0	0.0	5.81
CCDM 19236 - 3911 B	2.262	2.261	2.260	0.0	0.1	0.0	12.87
HD 122676	2.264	2.263	2.262	0.0	0.1	0.0	26.12
GJ 598	2.267	2.267	2.267	0.0	0.0	0.0	11.75
GJ 120.1 C	2.267	2.269	2.269	0.1	0.1	0.0	25.73
WD 0148+467	2.286	2.286	2.286	0.0	0.0	0.0	15.85
SIRIUS	2.299	2.299	2.299	0.0	0.0	0.0	2.64
HD 37574	2.290	2.305	2.233	0.7	2.5	3.1	62.00
HD 217107	2.300	2.313	2.323	0.6	1.0	0.4	19.72
HD 176687	2.299	2.314	2.333	0.7	1.5	0.8	27.31
BD -02 3986	1.804	2.316	3.102	28.4	72.0	33.9	58.48
HD 67852	2.932	2.341	1.229	20.2	58.1	47.5	118.34
HD 260564	2.341	2.341	2.340	0.0	0.0	0.0	34.42
ALGOL	2.481	2.381	2.666	4.0	7.5	12.0	28.46
GJ 54.1	2.429	2.429	2.429	0.0	0.0	0.0	3.72
LP 816 - 60	2.482	2.482	2.483	0.0	0.0	0.0	5.49
HD 50867	2.540	2.587	2.732	1.9	7.6	5.6	51.47
GJ 16	2.596	2.609	2.623	0.5	1.0	0.5	16.22
HD 34790	2.477	2.647	2.862	6.9	15.5	8.1	85.32
IRAS 17249+0416	2.670	2.664	2.658	0.2	0.4	0.2	50.08
HD 152912	2.860	2.700	2.466	5.6	13.8	8.7	139.28
GJ 768	2.702	2.702	2.702	0.0	0.0	0.0	5.14
GJ 903	2.791	2.791	2.792	0.0	0.0	0.0	13.79
HD 172748	2.801	2.806	2.823	0.2	0.8	0.6	57.34

**Table 4.** continued...

Star Name	Global	Local Missing Distance (pc)	Linear	(G-Lo)/G	(G-Li)/G (%)	(Lo-Li)/Lo	Dist (pc)
HD 142500	3.247	2.917	2.458	10.2	24.3	15.7	73.86
ROSS 614	2.929	2.929	2.929	0.0	0.0	0.0	4.12
GJ 1095	2.969	2.969	2.968	0.0	0.0	0.0	16.86
ROSS 882	3.053	3.052	3.052	0.0	0.0	0.0	5.93
HD 148317	3.268	3.132	2.903	4.2	11.2	7.3	79.87
HD 150689	3.145	3.145	3.146	0.0	0.0	0.0	14.37
GJ 169	3.189	3.189	3.188	0.0	0.0	0.0	11.47
GJ 279	3.196	3.196	3.197	0.0	0.0	0.0	25.70
GJ 628	3.208	3.208	3.208	0.0	0.0	0.0	4.26
GJ 687	3.213	3.213	3.213	0.0	0.0	0.0	4.53
GJ 231	3.249	3.249	3.249	0.0	0.0	0.0	10.15
HD 38382	3.271	3.271	3.273	0.0	0.1	0.1	25.54
GJ 71	3.271	3.271	3.271	0.0	0.0	0.0	3.65
GJ 1005	3.289	3.289	3.289	0.0	0.0	0.0	5.21
VAN MAANEN STAR	3.327	3.327	3.327	0.0	0.0	0.0	4.41
HD 28676	3.681	3.380	2.966	8.2	19.4	12.2	39.14
GJ 722	3.384	3.384	3.384	0.0	0.0	0.0	12.98
HD 233081	3.412	3.396	3.355	0.5	1.7	1.2	51.02
GJ 280 A	3.438	3.438	3.438	0.0	0.0	0.0	3.50
GJ 15A	3.467	3.469	3.469	0.1	0.1	0.0	3.57
GJ 678	3.503	3.503	3.503	0.0	0.0	0.0	16.45
GJ 725 B	3.515	3.515	3.515	0.0	0.0	0.0	3.52
HD 199881	4.355	3.527	3.106	19.0	28.7	11.9	80.45
HD 67523	3.563	3.563	3.564	0.0	0.0	0.0	19.23
GJ 725A	3.568	3.568	3.568	0.0	0.0	0.0	3.57
GJ 66	3.569	3.570	3.570	0.0	0.0	0.0	8.15
HD 168769	3.540	3.594	3.662	1.5	3.4	1.9	50.33
LUYTEN STAR	3.666	3.666	3.666	0.0	0.0	0.0	3.80
GJ 825	3.696	3.696	3.696	0.0	0.0	0.0	3.95
GJ 784	3.727	3.727	3.727	0.0	0.0	0.0	6.20
GJ 775	3.756	3.756	3.756	0.0	0.0	0.0	13.11
GJ 96	3.756	3.756	3.756	0.0	0.0	0.0	11.91
GJ 251	3.813	3.814	3.814	0.0	0.0	0.0	5.52
BD + 37 4901C	3.940	3.820	3.622	3.0	8.1	5.2	33.80
GJ 380	3.856	3.856	3.856	0.0	0.0	0.0	4.87
GJ 252	3.869	3.867	3.862	0.1	0.2	0.1	17.27
HD 122064	3.868	3.868	3.868	0.0	0.0	0.0	10.10
HD 168956	4.134	3.940	3.762	4.7	9.0	4.5	73.69
HD 53253	3.527	3.998	4.381	13.4	24.2	9.6	125.16
HD 146214	4.394	4.034	4.007	8.2	8.8	0.7	93.63
GJ 791.1A	4.021	4.053	4.103	0.8	2.0	1.2	30.27
GJ 268	4.066	4.066	4.066	0.0	0.0	0.0	6.36
HD 192869	3.979	4.080	4.072	2.5	2.3	0.2	111.98
HD 33959C	4.092	4.093	4.097	0.0	0.1	0.1	25.14
GJ 337.1	4.117	4.121	4.121	0.1	0.1	0.0	19.56
GJ 674	4.134	4.134	4.134	0.0	0.0	0.0	4.54
GJ 620.1A	4.153	4.155	4.158	0.0	0.1	0.1	12.87
GJ 103	4.180	4.180	4.180	0.0	0.0	0.0	11.51
GJ 851	4.205	4.203	4.203	0.0	0.0	0.0	11.44
HD 170296	4.256	4.278	4.280	0.5	0.6	0.0	89.37
GJ 222	4.380	4.380	4.380	0.0	0.0	0.0	8.66
GJ 752A	4.420	4.420	4.420	0.0	0.0	0.0	5.87
HD 218200	4.856	4.429	4.062	8.8	16.4	8.3	74.91
HD 207164	2.720	4.449	7.785	63.6	186.2	75.0	75.87
GJ 688	4.461	4.461	4.460	0.0	0.0	0.0	10.71
HD 58954	4.467	4.483	4.343	0.4	2.8	3.1	85.98
GJ 716	4.484	4.484	4.484	0.0	0.0	0.0	13.21
HD 32450	4.488	4.489	4.490	0.0	0.0	0.0	8.52
HD 162102	4.430	4.518	4.629	2.0	4.5	2.5	64.68
GJ 68	4.572	4.573	4.573	0.0	0.0	0.0	7.47
HD 71974	4.639	4.612	4.564	0.6	1.6	1.0	28.71
HD 163547	4.056	4.648	5.170	14.6	27.5	11.2	149.93

Table 5. continued...

Star Name	Global	Local Missing Distance (pc)	Linear	(G-Lo)/G	(G-Li)/G (%)	(Lo-Li)/Lo	Dist (pc)
ROSS 780	4.690	4.690	4.690	0.0	0.0	0.0	4.70
GJ 702	4.698	4.698	4.698	0.0	0.0	0.0	5.09
GJ 178	4.701	4.701	4.701	0.0	0.0	0.0	8.03
GJ 701	4.720	4.720	4.720	0.0	0.0	0.0	7.80
HD 72617	4.760	4.762	4.748	0.0	0.3	0.3	58.07
GJ 832	4.828	4.828	4.828	0.0	0.0	0.0	4.94
GJ 638	4.834	4.834	4.834	0.0	0.0	0.0	9.77
GJ 713	4.838	4.838	4.838	0.0	0.0	0.0	8.06
HD 239927	4.862	4.871	4.890	0.2	0.6	0.4	57.44
GJ 625	4.896	4.896	4.896	0.0	0.0	0.0	6.58
HD 67228	4.924	4.924	4.925	0.0	0.0	0.0	23.33
GJ 735	4.926	4.927	4.927	0.0	0.0	0.0	11.59
HD 75935	4.857	4.938	5.063	1.7	4.2	2.5	40.55
GJ 410	4.973	4.976	4.980	0.1	0.1	0.1	11.66
HD 39655	4.694	4.979	5.197	6.1	10.7	4.4	102.77
GJ 644	4.982	4.982	4.982	0.0	0.0	0.0	5.74
HD 120702	4.820	4.993	5.184	3.6	7.6	3.8	96.25

as a function of distance from the Sun. From this figure, we can appreciate that the three approximations give the same results up to distances of  $\sim 20$  pc. Even for times as short as 10 Myr, stars beyond  $\sim 20$  pc have significant differences in their trajectories for the different approximations, showing where the Galactic global potential becomes important. A local potential clearly does better than the straight line approximation. For completeness, we compare the results of our global model including arms and bar, with the simple global model of García-Sánchez et al. (2001), that includes spiral arms. In this case we find 73% of our trajectories differ by less than 3%, and for the rest (the most distant stars in general) the error is 6-26%. Differences among stellar orbits due to the use of different models for the Galactic field or the simple stright-line approximation, become more important the farther the star is, making even more uncertain approaches to the Sun in its path for the Milky Way disc. Deeper and precise observations as the ones are coming in the near future, will improve dramatically our knowledge on this field, and will make more important to use better models of the Galaxy.

### 3.1.4 The Solar System Under the Influence of Stellar Encounters at the Solar Neighborhood

The overarching goal of our work in this set of papers will be to address quantitatively the effect of stellar encounters in different Galactic environments. In this section, we start with the nascent Solar System and the current stellar environment. For this purpose, we model a simple planetary system (or disk) with 1000 particles distributed from 1 to 100 AU. We study this system for a total integration time of 10,000 yr, which is much longer than the typical encounter timescale.

We chose the parameters of Gliese 710 ( $0.6 M_{\odot}$ , approximation velocity  $13.9 \text{ km s}^{-1}$ ) and an estimated closest approach less than 0.3 pc, 1.4 Myr in the future, to calculate its effects in the Solar System. The effect of this specific star

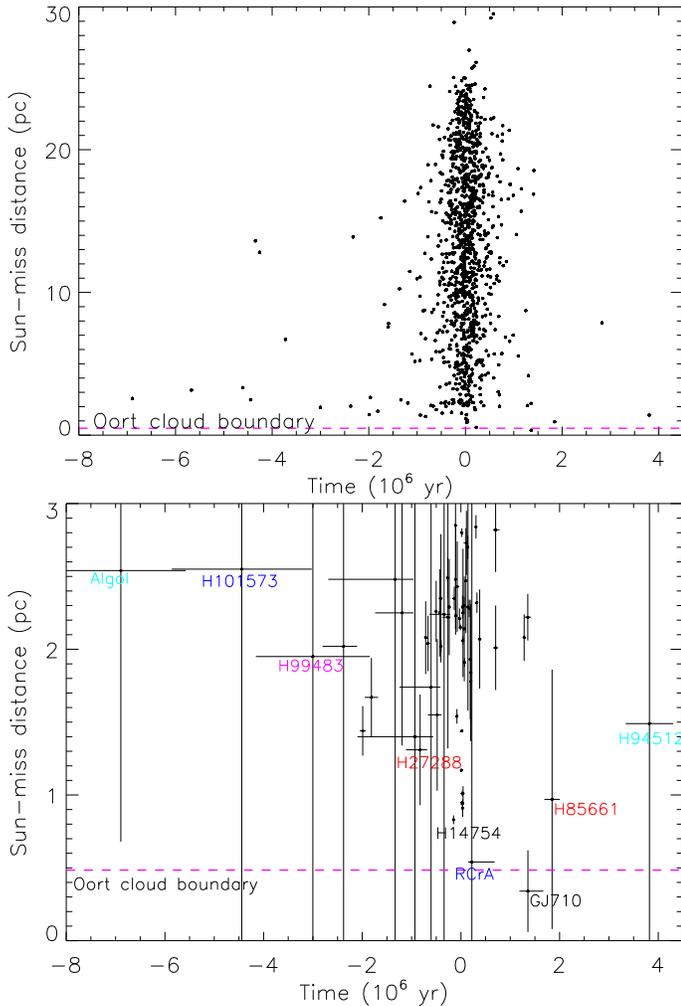
will be interesting for the Oort cloud, but we find that Gliese 710's effect is negligible for a 100 AU planetary system.

We ran a grid of simulations for a flyby star mass of  $1 M_{\odot}$  and closest approaches between 100 AU and 1000 AU, in steps of 50 AU. Disk disruption starts to be significant for approaches within about 200 AU. We also find a good agreement to the analytical result of Hall, Clarke & Pringle (1996) where they find that a disk is affected to  $\sim 1/3$  the closest approach distance. With these parameters and the velocity of GJ 710 ( $13.9 \text{ km s}^{-1}$ ), we calculate the gravitational effect on a planetary disc and we present it in Figure 6. The figure shows four panels, representing the orbits of particles in x-y (left frames) and x-z (right frames) planes, before (upper frames) and after (lower frames). Figure 7 shows the resultant orbital characteristics of the disc after the encounter, eccentricity of particles (upper left frame), inclination (upper right frame), pericenter distance (lower left frame) and apocenter distance (lower right frame), all plotted versus the semimajor axis. The effect on the disc is slight but clear starting at 40 AU, where particles reach an eccentricity up to 0.1.

### 3.1.5 Effect of Gliese 710 on the Oort Cloud

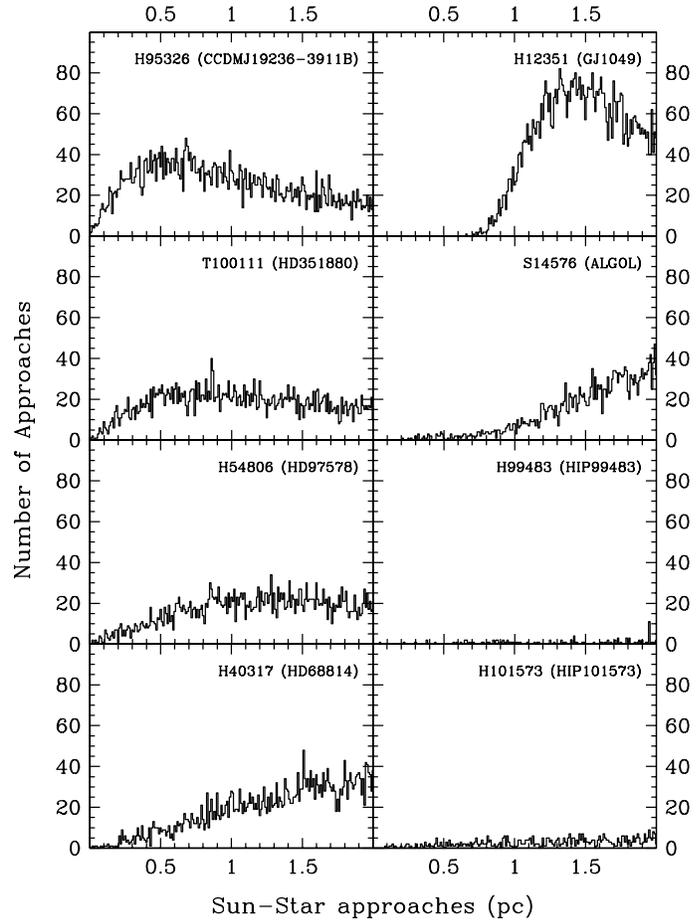
The solar system's disk and Kuiper belt are surrounded by the Oort cloud, which contains  $1 \times 10^{11}$  to  $5 \times 10^{12}$  cometary nuclei with a total mass  $\sim 1$  to  $50 M_{\oplus}$  (Stern 2003). Comets in the Oort cloud evolve dynamically under the influence of external perturbers such as random stellar passages. Close or penetrating passages of stars through the Oort cloud can deflect comets toward the inner planetary region (Hills 1981; Weissman 1996; Kirsh et al. 2009; Brassier et al. 2006, 2007, 2008).

One important issue in the scenario of the flyby star is the effect it would have on the loosely bound outer Oort cloud. Small gravitational perturbations could have severe effects on the cloud due to the direct effect on the Sun. A simple calculation using the impulse approximation predicts

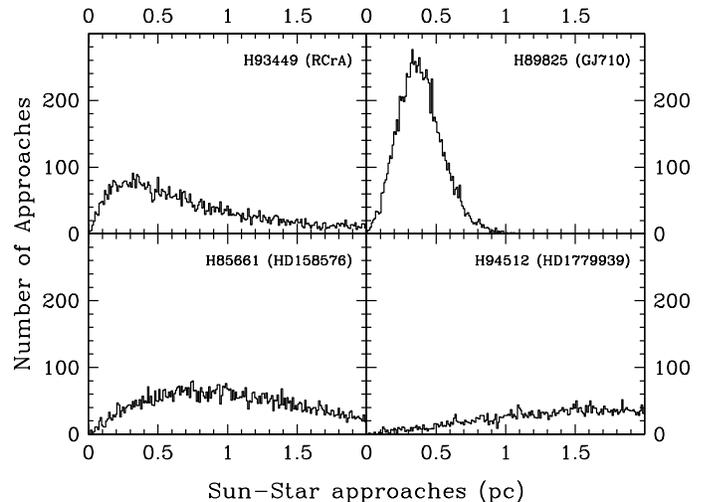


**Figure 2.** Miss distance vs. time of stellar approaches for those stars in the total sample of 1167 stars within 30 pc as maximum approach distance to the Sun. The Oort cloud boundary at  $\sim 0.48$  pc is plotted as reference. The lower frame is a zoom that shows a subsample with the closest approach distances (less than 3 pc) to the Sun (see Tables 1 and 2) and their computed uncertainties. Some star’s names are illustrated in the figure.

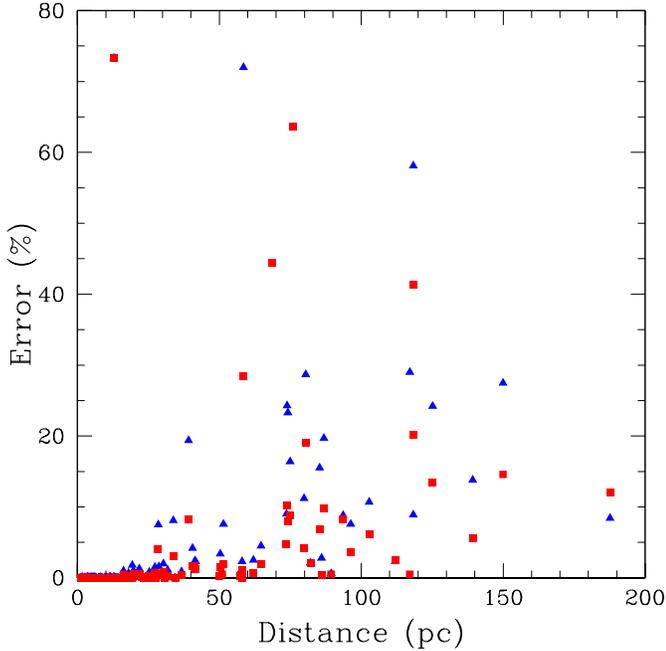
that during a stellar encounter the velocity of the Sun would change by  $\delta v \approx \frac{2G M_*}{q_* v_\infty}$ . With the values of Gliese 710, for the stellar mass,  $M_* = 0.6 M_\odot$ , the closest distance of the encounter,  $q_* = 0.34$  pc ( $\sim 70,000$  AU) and the velocity at infinity,  $v_\infty = 13.4$  km s $^{-1}$ . This would result in  $5.8 \times 10^{-4}$  km s $^{-1}$  change in the solar velocity, a negligible value compared with the typical Oort cloud speed of 0.2 km s $^{-1}$  or with its escape velocity at the boundary ( $\sim 0.1$  km s $^{-1}$ ), so Gliese 710 has no ability to strip the Oort cloud from the Sun. However, it has the potential to send a comet flux toward the inner Solar system. This will cause the appearance of a new comet per year and the net increase of the risk of astronomic impact will not be detectable (García-Sánchez et al. 2001).



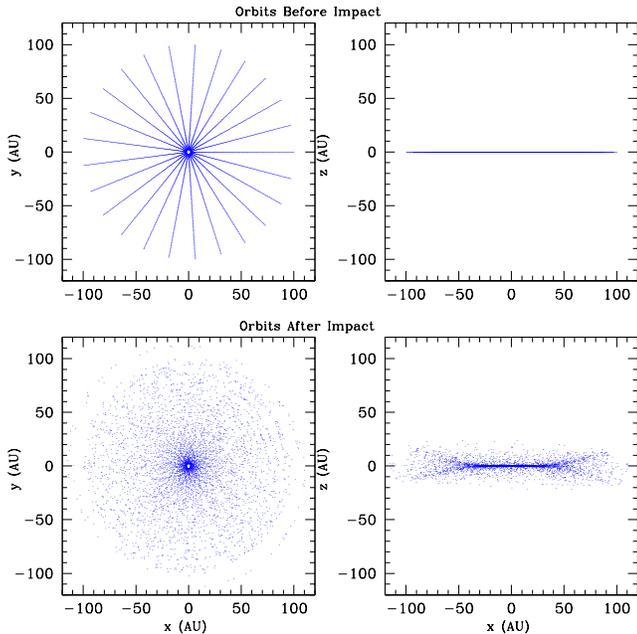
**Figure 3.** Histograms computed on eight stars from our sample to the past with large enough observational error bars that may produce approach distances between the Sun and the star below zero. All plots show the Sun-star approaches in pc from a set of 10,000 normal deviation random initial conditions within observational error bars.



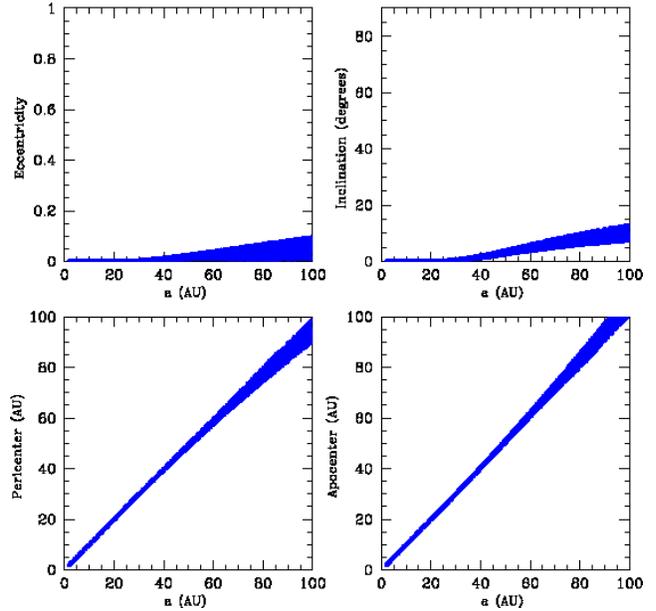
**Figure 4.** Same as Figure 3 but for four stars to the future.



**Figure 5.** Relative error between the global and the straight line approximations to the Galactic potential (blue triangles) and between the global and local approximations (red squares).



**Figure 6.** Effect of a stellar flyby mass of  $1 M_{\odot}$ , closest approach of 300 AU and velocity of  $13.9 \text{ km s}^{-1}$  on a disc of particles. Left frames: orbits of the particles in x-y cuts, right frames: cut in the x-z plane. Upper frames: particles before the encounter, lower frames: particles after the encounter



**Figure 7.** Resultant orbital characteristics on the disc of Figure 6 after the encounter. Upper left frame: eccentricity, upper right frame: inclination, lower left frame: pericenter distance, and lower right frame: apocenter distance. All versus the semimajor axis

#### 4 SECOND GALACTIC REGION: THE BIRTH CLOUD OF THE SUN

In star formation regions such as Sun’s putative birth location inside a dense cluster, stellar densities are high enough to produce stellar encounters within 300 AU before the dissolution of the stellar cluster (Laughlin & Adams 1998; Adams 2010). Also, in this environment, encounters are stronger owing to lower typical velocity dispersions between 1 and  $3 \text{ km s}^{-1}$ . Let us consider, for example, a planetary disc around a given star on a crowded stellar environment. Let us assume solar mass stars for generality. The star-disc system will experience encounters with other systems with an interaction rate that can be written as  $\gamma = \langle n\sigma v \rangle = \langle n \rangle \langle \sigma \rangle \langle v \rangle$ , where  $n$  is the number density of stars in the cluster,  $\sigma$  is the cross section for interaction, and  $v$  is the relative velocity (velocity dispersion in this case). For a rough estimate of the interaction rate among stars, we use the typical values for open clusters. In their work, Laughlin & Adams (1998), calculate the interaction rate for the Trapezium cluster, using a central density of  $n_0 \approx 5 \times 10^4 \text{ pc}^{-3}$  and a velocity dispersion of a few  $\text{km s}^{-1}$ , and an interaction cross section of  $100^2 \text{ AU}^2$ , obtaining a rate of interaction of one encounter every 40 million years. Considering open clusters live for at least 100 million years, it is expected these interactions are significant in those environments. In the case of the Sun’s birth cloud several observables, such as the inclinations of Uranus and Neptune (which are sensitive to stellar interactions) place densities at similar amounts of typical open clusters, this is  $\sim 3 \times 10^4 \text{ M}_{\odot} \text{ pc}^3$  (Gaidos 1995). Furthermore, the orbit of (90377) Sedna supports the idea that the Solar system was born in a stellar cluster with a non-negligible density

(Brasser et al. 2006; Kenyon & Bromley 2004; Morbidelli & Levison 2004).

In this section, we present experiments to reproduce some of the Kuiper belt bodies orbital characteristics, such as eccentricities and inclinations. We have chosen four close-approach distances (200, 150, 100 and 50 AU), coupled with three initial approach velocities (1, 2 and 3 km s<sup>-1</sup>) for a stellar flyby interacting with a 100 AU disc of particles. The mass of the perturbing star is 1 M<sub>⊙</sub>, and angles involved in the geometry of stellar encounter have the general values,  $\theta=45^\circ$  (angle between the plane of the stellar trajectory and the plane of the disc) and  $\alpha=45^\circ$  (angle between the axis of the stellar trajectory plane and the axis of the disc).

In our initial conditions, test particles have  $e = i = 0$ . In Figure 8 we show our results of perturbations produced in the disk of particles by a flyby. The columns represent different values of the initial approach velocity (1, 2 and 3 km s<sup>-1</sup>) and rows represent different values of closest approach distances (50,100,150 and 200 AU). Each panel shows the positions of the test particles in the x-z plane after the encounter with the flyby star.

In Figures 9 and 10 we show the resultant eccentricities and inclinations, plotted versus semi-major axis. Columns represent different distances of closest approach, within 200 AU, and rows indicate different values of velocity dispersion, within 3 km s<sup>-1</sup>. As a reference, in all plots, we include the known Kuiper belt objects, including the resonant and classic objects (pink triangles), and the scattered objects, including also Centaurs at radii less than 30 AU (green crosses).

Our best fit happens at a distance of maximum approach between 100 and 150 AU and a velocity of  $\sim 1$  km s<sup>-1</sup>, with an unaltered inner planetary system (as far as 30 AU) where the most of particles between 30 and 50 AU have up to  $e = 0.2$  and  $i < 20^\circ$ . In this simulation we also obtain some dispersed objects with high eccentricities and low semimajor axis as many of the objects seen in the planetary system. From 42 to approximately 48 AU there are objects with eccentricities up to 0.4. We are able to match these objects using encounters with a velocity at infinity of 1 km s<sup>-1</sup> and a close-approach distance of 100 AU. This results in eccentricities from 0-0.1, on semimajor axis interval 0-40 AU, and eccentricities from 0 to 1 on semimajor axis in the interval 40-65 AU and from 65 until 100 AU the most of particles are dynamically evaporated, this means that objects with low eccentricity and large semimajor axis (larger than 50 or 60 AU) are scarce in these experiments as it is the case in the Kuiper belt.

## 5 CONCLUSIONS

This paper is the first step in an extensive study of the gravitational effect of a flyby star on a 100 AU disc, representing debris discs (and/or planetary systems) given the particular conditions of several environments of the Galaxy. In this work, we analyzed the effect of the Solar neighborhood and the birth cloud of the Sun. In the current Solar neighborhood, we find that Gliese 710 will be the closest star to the Sun 1.36 Myr in the future with a minimum distance of 0.34 pc ( $\sim 70,000$  AU). This stellar encounter will lead to direct interactions between the star and the Oort cloud. This star

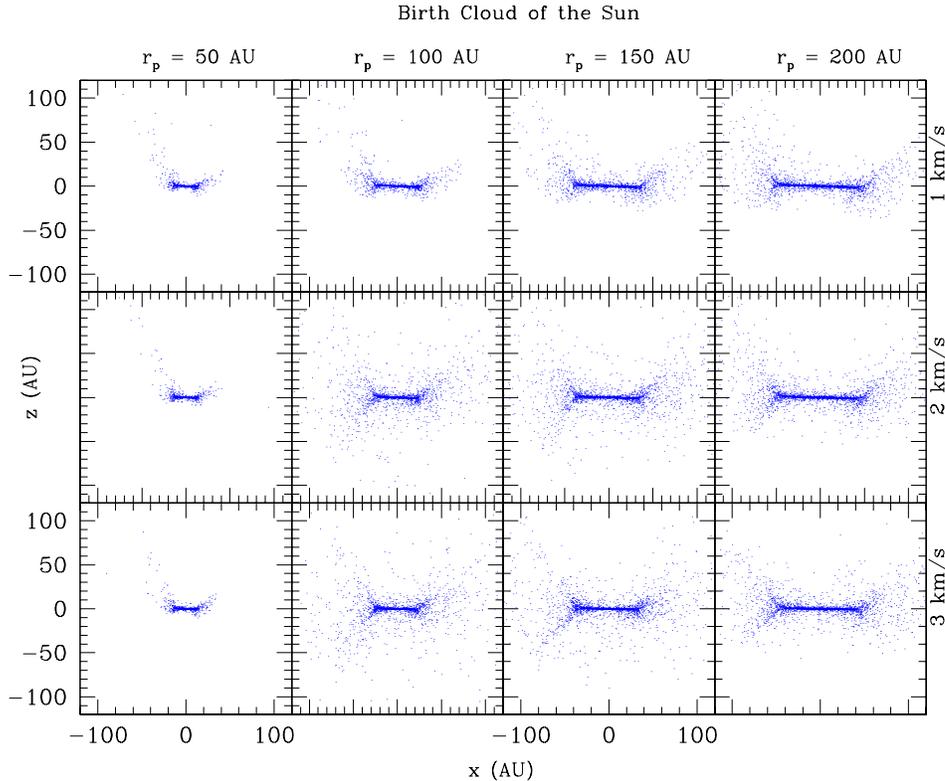
has a mass of 0.6 M<sub>⊙</sub> currently located at 19.3 pc from the Sun. We calculated its velocity at the point of minimum distance with the Sun as 14 km s<sup>-1</sup>. The effect generated by Gliese 710 on a 100 AU planetary system will be negligible. The minimum distance for which the effect caused by Gliese 710 would start to be noticeable is  $\sim 300$  AU, producing a slight heating of the outer parts of the 100 AU planetary disc, with eccentricities up to 0.1 on semimajor axes between 60-100 AU and inclinations up to 10°. With a simple impulse approximation we find that Gliese 710 will not even have an important effect on the global structure of the Oort Cloud.

For stars of the Solar neighborhood, with 6D known information, we constructed orbits in order to compare the use of a simple straight line approximation, a local approximation to the Galactic potential, an axisymmetric potential and a full Galactic potential (including spiral arms and bar observationally motivated). Even for times as short as 10 Myr, stars beyond  $\sim 20$  pc have significant differences in their trajectories for the different approximations, showing where the Galactic global potential becomes important. A local potential clearly does better than the straight line approximation. Comparing the global potential model (including spiral arms and bar), with the simpler model of García-Sánchez et al. (2001) (that includes spiral arms), we find that 73% of our trajectories differ by less than 3%, and for the rest (the most distant stars in general) the error is 6-26%. Differences among stellar orbits due to the use of different models for the Galactic potential field or the simple straight line approximation, become more important the farther the star is, making even more uncertain approaches to the Sun in its path for the Milky Way disc. To compute precise stellar orbits of the Solar neighborhood, taking advantage of the new generation of data produced by large surveys, will require better models of the Milky Way Galaxy instead of simpler approximations.

Finally, regarding the birth cloud of the Sun, we produced several experiments to reproduce the orbital parameters of the Kuiper belt objects. We know that at 42-48 AU there are objects with eccentricities up to 0.4. We are able to approximate these objects using encounters with a velocity at infinity of 1 km s<sup>-1</sup> and a close-approach distance of 100 AU. This results in eccentricities from 0 - 0.1, on semimajor axis interval 0-40 AU, and eccentricities from 0 to 1 on semimajor axis in the interval 40-65 AU and from 65 until 100 AU the most of particles are dynamically evaporated.

## ACKNOWLEDGMENTS

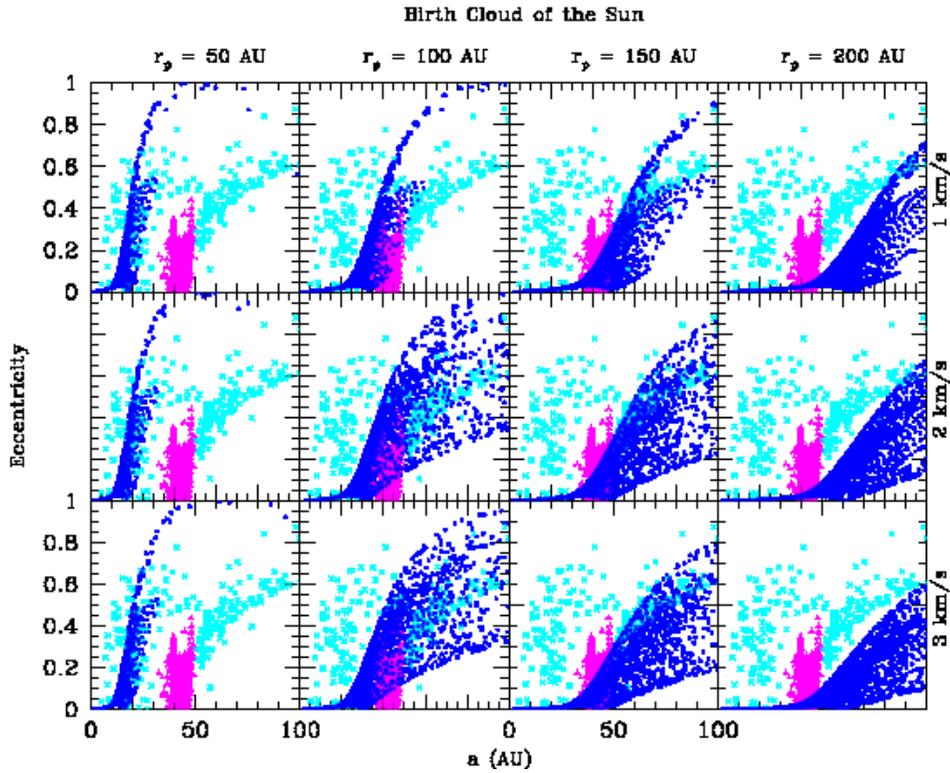
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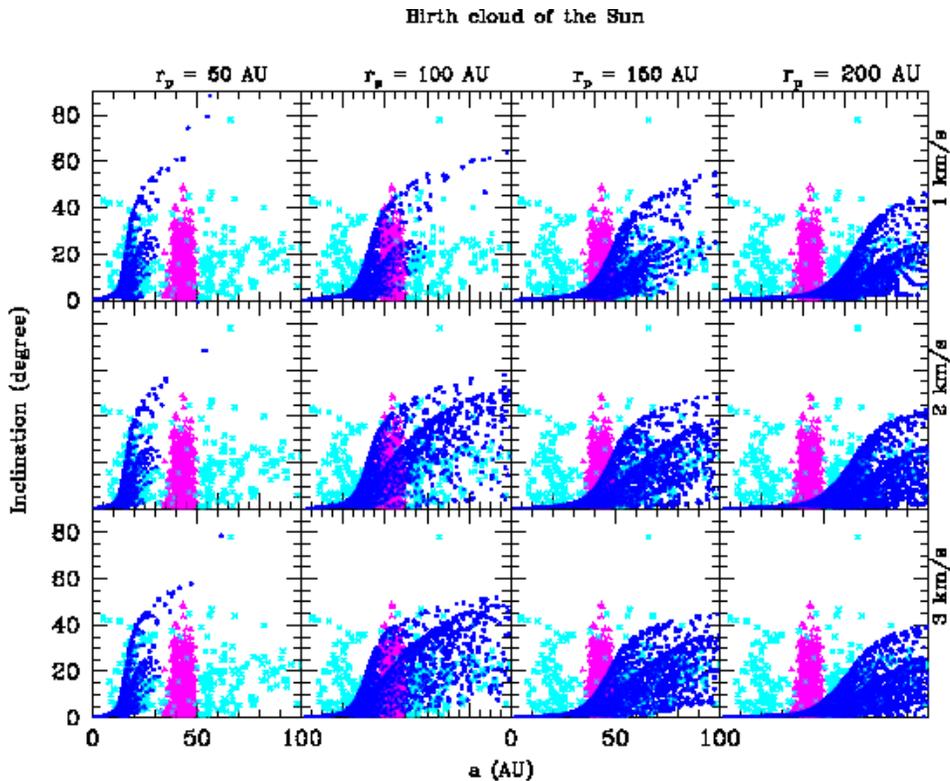
**Figure 8.** Final disc structure after a stellar encounter simulating a Galactic environment of star formation. Rows represent values of velocity dispersion within  $3 \text{ km s}^{-1}$  and columns represent different impact parameters within 200 AU

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**Figure 9.** Resultant eccentricities. Rows indicate different values of velocity dispersion and columns represent different distances of closest approach. Our best fit happens for a stellar encounter with minimum distance of 100-150 AU and a velocity of  $\sim 1 \text{ km s}^{-1}$ . Resonant objects and classic Kuiper belt objects are included (pink triangles), and also scattered objects and Centaurs at radii less than 30 AU (cyan crosses)



**Figure 10.** Same as Figure 9 but for inclinations

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